

Broad-band radio polarimetry of disc galaxies and AGN with KAT-7

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ABSTRACT

We report broad-band (1.2–1.9 GHz) radio continuum observations at arcminute resolutions of two nearby disc galaxies, NGC 1808 and NGC 1097, and four active galactic nuclei (AGN)-powered radio sources: PKS B1934–638, PKS B0407–658, J0240–231, and J0538–440. We use rotation measure synthesis to analyse their Faraday complexity. Observations were made with the seven-dish Karoo Array Telescope (KAT-7) radio telescope array, in South Africa. The AGN-powered sources fall into two ‘Faraday’ categories – simple and complex. The most polarized sources, J0538–440 and J0240–231, are found to have complex Faraday spectra that can be time variable (J0538–440 case) and also indicative of complex Faraday emitting and rotating components along the line of sight. PKS B0407–658 shows a simple Faraday spectrum, while PKS B1934–638 is undetected in polarization. The disc galaxies are classified as complex, albeit at low signal-to-noise ratio. This may indicate depolarization due to turbulence of the magnetized plasma in the bar and circumnuclear regions and/or frequency-dependent depolarization at L band.

Key words: polarization – galaxies: magnetic fields – radio continuum: galaxies.

1 INTRODUCTION

How magnetism fits into the picture of galaxy evolution over cosmic time is a subject of debate among astronomers. The link between magnetic fields and galaxy evolution is theorized (e.g. Lou & Fan 1998) but observational constraints have been lacking, due to technical limitations (e.g. Chamandy, Shukurov & Taylor 2016). Understanding the evolution of magnetic fields and their relation to galaxy evolution is one of the key science objectives of next generation radio telescopes such as the Square Kilometre Array (SKA; Johnston-Hollitt et al. 2015; Heald, Mao & Vacca 2020). The polarization of radio emission is a powerful probe of cosmic magnetic fields.

Some observational studies in the radio (Eichendorf & Reinhardt 1979; Beck et al. 2002, 2012; Beck 2004; Farnes 2014) have provided insights into the workings of magnetism in galaxies. Highly sensitive and resolved imaging observations of nearby disc galaxies (Beck et al. 2005) show large-scale magnetic fields related to spiral structure. For more distant galaxies that remain unresolved by current radio imaging telescopes, we must rely on the frequency dependence of polarization, instead of direct imaging, as our primary probe of magnetic properties (Stil et al. 2009a). However, the relatively narrow instantaneous bandwidths typical of most past studies of radio polarization limits both the resolution of Faraday spectral features and the ability to characterize the frequency-dependent effects of internal depolarization.

Recent studies have begun to probe polarized radio emission over broader bands. For example, O’Sullivan et al. (2012) presented a broad-band spectropolarimetric study of four radio bright and highly polarized active galactic nuclei (AGN) sources, for which Faraday rotation features could be resolved. Anderson et al. (2015) report a wide band study of polarimetry in more than 500 sources, Anderson, Gaensler & Feain (2016) observed 36 strongly polarized sources over a very wide band from ~ 1 to 10 GHz, and find Faraday complexity in most of their sources, asserting that complexity is thus common among strong polarized sources. A more recent study, Schnitzeler et al. (2019), provides a wide-band (1.3–3.1 GHz) polarimetry survey of unresolved (angular resolution $\sim 2 \times 1$ arcmin²) southern sources. The authors argue in favour of the Galactic foreground being a dominant component in terms of Faraday rotation of background sources, implying simple rather than complex intrinsic distributions of Faraday rotation.

In this paper, we extend the studies of broad-band polarimetry of extragalactic radio sources by studying a sample of six radio sources, including two disc galaxies and four AGN-dominated sources. The study was carried out with observations from the seven-dish Karoo Array Telescope (KAT-7), which was constructed as an engineering prototype array at the South African SKA site as part of the technical development of the MeerKAT SKA precursor telescope. The 64-dish MeerKAT array will form part of the SKA mid-frequency component. The KAT-7 wavelength coverage from 1.2 to 1.9 GHz provides us with a good resolution of the Faraday spectrum of polarized emission (see Section 3.2). The rotation measure synthesis technique (RM Synthesis), first described by Burn (1966), is employed to explore the intrinsic polarization properties of the sources.

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The spatial resolution limitations of KAT-7 at low redshift are comparable to those of the more sensitive MeerKAT (and the future full SKA array) at high redshift. MeerKAT has the sensitivity to detect polarization in faint distant objects but does not resolve them, so we explore the peak flux properties of nearby objects to probe the spectropolarimetric behaviour in a fashion that can be extended to more distant and fainter objects with much more powerful instruments. This paper is structured as follows: Section 2 discusses our observations and data reductions; Section 3 presents the results of our polarimetric analysis; we discuss our findings in Section 4; and we summarize, discuss, and also make conclusions based on this work in Section 5.

2 OBSERVATIONS AND DATA REDUCTION

2.1 Sample selection

We selected our sample of disc galaxies from the source list of the MeerKAT H I Observations of Nearby Galactic Objects: Observing Southern Emitters survey¹ (MHONGOOSE; de Blok et al. 2018). The survey will secure very high sensitivity atomic hydrogen images of a sample of nearby galaxies. A further criterion was that our sources be sufficiently radio bright, with a reported integrated total intensity greater than 200 mJy at L band. The flux limit was chosen so as to have sources with sufficient flux density to be detectable by KAT-7 at a few per cent polarization. This resulted in a sample of two disc galaxies. We also observed four AGN-dominated sources, with three of them serving as calibrators for the remaining targets. The two disc galaxies both have unique features in the context of the overall disc galaxy population, but the analysis done here will shed some light on the polarization resulting from the large-scale magnetic field in disc galaxies that are unresolved.

2.1.1 Disc galaxies

NGC 1097 is identified as a type 1 Seyfert (Sy1) galaxy (Storchi-Bergmann et al. 1997) with a circumnuclear starburst ring (Gerin, Combes & Nakai 1988). It is a barred spiral, SBb, with its radio emission dominated by the circumnuclear ring and nucleus (Beck et al. 2005). The magnetic properties of NGC 1097 have been studied using narrow-band (~ 50 MHz) high-resolution radio data (2–15 arcsec resolutions) at 4.8 and 8.4 GHz, with the National Radio Astronomy Observatory (NRAO) Very Large Array (VLA; Beck et al. 2002, 2005). The galactic magnetic field was observed to control the interstellar medium (ISM) gas flow at kpc scales and found to have a regular component that follows the shape of the spiral arms (Beck et al. 2005). The galaxy has a redshift of 0.0042.

NGC 1808 is identified as a type 2 Seyfert (Sy2) galaxy (Tacconi-Garman, Sternberg & Eckart 1996; Galliano et al. 2008) with a low-luminosity AGN (Véron-Cetty & Véron 1985; Dahlem et al. 1990; González-Martín et al. 2013; Esquej et al. 2014). It was classified as a peculiar Sbc galaxy (Véron-Cetty & Véron 1985). It also has a starburst region in the inner ~ 750 pc dominating the infrared emission (Sengupta & Balasubramanyam 2006). From observations at 6 and 20 cm wavelengths with the VLA, Dahlem et al. (1990) found linear polarization degrees of up to 30 per cent and a steep ($\alpha = -0.88$) radio spectrum from outside the central region of the galaxy. The galaxy has a redshift of 0.0033.

2.1.2 AGN

J0240–231 has been identified to be a quasi-stellar object (QSO) and also a gigahertz-peaked spectrum (GPS) source (Kuehr et al. 1981; O’Dea 1998; Véron-Cetty & Véron 2006). It has a redshift of $z = 2.22$ (Hewitt & Burbidge 1987). The source has been observed to show significant linear polarization at 4.8 GHz where $p = 3.47$ per cent (Edwards & Tingay 2004). Very long baseline interferometry (VLBI) studies show structure resembling two lobes that are separated by 12 mas (Dallacasa et al. 1998).

J0538–440 is reported to be a BL Lac, GPS object, and also a variable γ -ray source located at redshift 0.896 (Impey & Tapia 1988; Romero, Cellone & Combi 2000; Tornikoski et al. 2001; Romero et al. 2002; Andruchow, Romero & Cellone 2005; Tornainen et al. 2005). It has also been identified as a potential gravitational lens (Surpi, Romero & Vucetich 1996). This source presents a time variable total intensity spectrum, at cm and mm wavelengths, as can be seen in calibrator data bases such as the Australia Telescope Compact Array (ATCA) calibrator data base.² It has also been observed to be highly polarized and variable in the optical (Impey & Tapia 1988; Romero et al. 2000, 2002; Andruchow et al. 2005).

PKS B1934–638 is a quasar and GPS source at redshift 0.18, while *PKS B0407–658* is a quasar and possible compact steep-spectrum (CSS) source (O’Dea, Baum & Stanghellini 1991; Labiano et al. 2007) at a redshift of 0.96. VLBI imaging of PKS B1934–638 shows that it has a double-lobed structure (Tzioumis et al. 1989). Henceforth, we use the name ‘PKS 1934–638’ for PKS B1934–638 in this work as this name is most commonly used for the source.

2.2 Observations

KAT-7 is a synthesis array telescope located adjacent to the MeerKAT and SKA phase 1 central site in the Karoo plateau of South Africa. It was built as an engineering testbed preceding the 64-antenna SKA mid-frequency precursor array, MeerKAT. KAT-7 consists of seven 12-m, centre-fed parabolic antennas in a fixed configuration with baselines ranging from 26 to 185 m (Carignan et al. 2013). It has a maximum instantaneous bandwidth of 256 MHz within the radio frequency (RF) range 1200–1950 MHz. Two components (X and Y) of linear polarization are detected with prime-focus feeds on each antenna. We observed in wide-band, full-polarization mode, providing raw data consisting of four polarization correlation products sampling a 256 MHz band in 1024 spectral channels, each of width 390.625 kHz. Our observations covered the entire 750 MHz RF band using three separate observing runs for each disc galaxy target with the 256 MHz centred near 1350, 1600, and 1850 MHz (see Table 1).

Each observation consisted of 12 h tracks during which on-source observations of 4 min of either NGC 1097 or NGC 1808 were alternated with 2 min scans of a phase calibrator, either J0240–231 (for NGC 1097) or J0538–440 (for NGC 1808). Flux and bandpass were calibrated with several 3 min scans on PKS 1934–638 and PKS B0407–658 depending on which calibrator was visible. The cadence on these flux calibrator scans was approximately an hour. The 12 h tracks provided a good range of parallactic angle variation for polarization calibration.

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²https://www.narrabri.atnf.csiro.au/calibrators/calibrator_database_viewcal.html?source=0537-441

Table 1. Summary of the KAT-7 observations. The table lists the radio sources (some aliases in brackets) observed during each run, their J2000 coordinates, the observation dates (and times given in UTC), central frequency (ν_c), number of antennas available (N_{ant}), and the total recovered bandwidth ($\Delta\nu$).

Source (aliases)	RA (J2000) (^h : ^m : ^s)	Dec. (J2000) ([°] : ['] : ^{''})	Date (UTC time)	ν_c (GHz)	N_{ant}	$\Delta\nu$ (MHz)
NGC 1808	05:07:42.3000	−37:30:47.0000	29 July 2016 (01:07:20.1–07:12:30.4)	1.894	6	152
PKS 1934–638 (J1939–6342)	19:39:25.0267	−63:42:45.6255				
J0538–440 (J0538–4405, 0537–441)	05:38:50.1800	−44:05:10.3000				
PKS B0407–658 (J0408–6545)	04:08:20.3788	−65:45:09.0806				
–	–	–	–	–	–	–
NGC 1808			31 July 2016 (01:00:19.5–07:03:29.8)	1.394	6	148
PKS 1934–638						
J0538–440						
PKS B0407–658						
–	–	–	–	–	–	–
NGC 1097	02:46:19.0000	−30:16:30.0000	09 June 2015 (02:32:11.9–08:34:42.2)	1.850	7	223
PKS 1934–638						
J0240–231 (J0240–2309, 0237–233)	02:40:08.1400	−23:09:16.0000				
PKS B0407–658						
–	–	–	–	–	–	–
NGC 1097			15 June 2015 (02:30:22.9–08:35:13.2)	1.350	7	211
PKS 1934–638						
J0240–231						
PKS B0407–658						

2.3 Data reduction

Processing of the raw visibility data was carried out using the Common Astronomy Software Applications (CASA; McMullin et al. 2007) package. After removal of the channels at the upper and lower ends of the band due to roll-off of the bandpass shape, the visibility data were inspected visually and manually flagged to remove radio frequency interference (RFI) signals. The mid-band, centred at 1600 MHz, was heavily populated with RFI. Analysis in this band was severely compromised as a result and we did not include the mid-band in our subsequent analysis. The objects observed, recovered band pass, and number of operating antennas for each observing run are listed in Table 1.

PKS 1934–638 was used for absolute flux calibration, and to measure the complex bandpass shape and cross-hand delay for all observations. The sources J0538–440 and J0240–231 were used as secondary calibrators to track time-dependent amplitude and phase calibration over the course of each observation and for on-axis polarization leakage calibration given the large parallactic angle covered. The time-dependent total intensity gain is coupled to any detectable polarization signal from the astronomical sources and telescope. The broad parallactic angle coverage of the observations of the secondary calibrators allows us to solve for the polarization of the calibrators from the initial antenna-based time-dependent gain solutions and the frequency-dependent XY-phase calibration. The source polarization solutions are then used to correct the time-dependent gains and to solve for frequency-dependent polarization leakage. CASA polarization calibration tasks do not account for rotation measure (RM) across a wide band and thus the polarization solution from the gains is the average across the band. This is not ideal for a non-zero RM. This effect is not mitigated during data reduction in this work and thus may affect polarization measurements. We assume that the effect will not be significant at the bandwidths of our data, ~ 200 MHz, and that none of the targets have large RM values. For example, Taylor, Stil & Sunstrum (2009) report an RM of 14.5 ± 2.2 rad m^{-2} with polarization degree of 0.8 ± 0.1 per cent at 1.4 GHz for J0240–231.

We do not observe a source with known absolute polarization position angle. Therefore, while the relative polarization position

angles (with respect to frequency) are calibrated, the absolute polarization position angles for the polarization solutions are not known. The instrumental polarization is accounted for by solving for leakages using a source with a large parallactic angle coverage (e.g. Jagannathan et al. 2017).

After applying the calibration solutions, Stokes I , Q , and U image cubes were made for each source. The imaging and deconvolution were done using the Clarke–Stokes algorithm implemented in CASA’s TCLEAN task. We averaged the visibility channels into 3.9 MHz channels to construct the image cubes that resulted in the maximum detectable Faraday depth decreasing by an order of magnitude (see Section 3.2 and Table 4). The typical synthesized beam at the low band is ~ 4.9 arcmin and ~ 3.6 arcmin at the high band. The observed angular size of our sources is consistent with unresolved sources.

3 RESULTS

3.1 Spectropolarimetry

We extracted spectra in Stokes I , Q , and U for each source at the peak total intensity in each cube channel. The errors for each of the peak intensities are estimated by taking the standard deviation of intensities at source-free/off-source regions in each cube channel. We then propagate these errors for derived quantities, such as linear polarization, calculated from I , Q , and U . The peak intensity values are measures of the integrated intensity, as these sources are unresolved.

Per-channel Stokes I intensities for each source are shown as blue dots in each panel of Fig. 1. For comparison we also show flux densities for each source over a broader frequency range obtained from the NASA/IPAC Extragalactic Database (NED)³. There is very good agreement with previously reported values for PKS 1934–638, NGC 1808, and PKS B0407–658. The spectrum, for NGC 1097,

³The NASA/IPAC Extragalactic Database (NED) is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

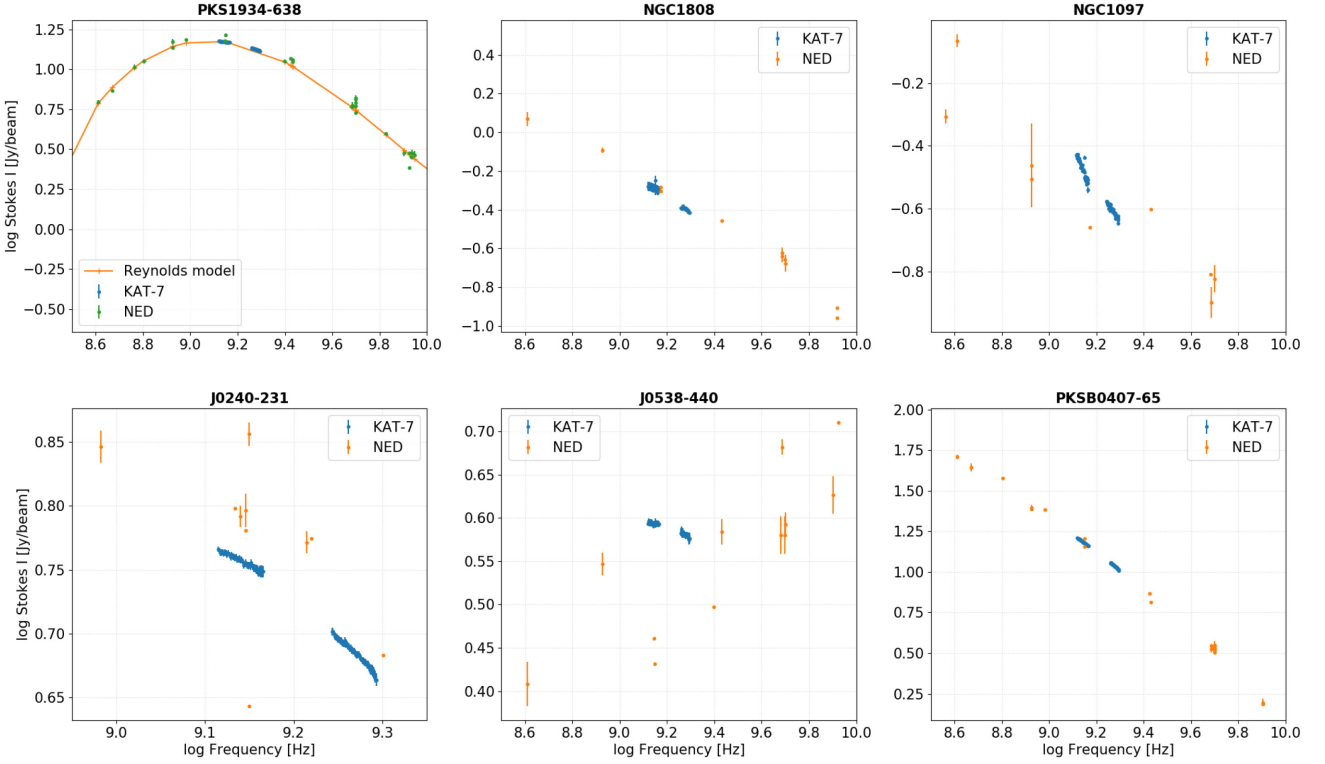


Figure 1. Total intensity across KAT-7 band (blue points), some corresponding literature values (from NED, <https://ned.ipac.caltech.edu/>) at frequencies near the KAT-7 band. Top panels: the primary flux calibrator, left: PKS 1934–638, with the Reynolds (1994) model plotted as a red curve; middle: NGC 1808; right: NGC 1097. Bottom panels: the polarization calibrators, left: J0240–231; middle: J0538–440; right: PKS B0407–658.

while consistently steep, is very poorly defined by previous data. However, our data fall within the general trend of past observations. The data for J0538–440 suggest variability. Our measured spectrum differs significantly from the literature values. The discrepancy is brought about most likely by time variability in the J0538–440 spectrum as can be seen in the ATCA data base. The source has a total intensity spectrum that is flat near ~ 1 GHz with a standard deviation of ~ 1.0 Jy in the region, while showing a bump towards, and also reaching a peak near, ~ 10 GHz. The results for the flat spectrum source J0240–231 also indicate variability. The total flux of J0240–231 in the 1.36–2.0 GHz frequency range has been reported to vary over the range 4.4–7.1 Jy (Kuehr et al. 1981; Condon et al. 1998; Stanghellini et al. 1998, 2005; Tingay et al. 2003).

We derive the bias-corrected (Simmons & Stewart 1985) linear polarized flux density as

$$P_o = \sqrt{P^2 - \sigma_{QU}^2}, \quad (1)$$

where $P^2 = Q^2 + U^2$ and σ_{QU} is the noise in the peak Q and U values derived as the rms of the amplitudes of off-source positions in each channel. For channel-averaged quantities $\langle Q \rangle$ and $\langle U \rangle$, the noise is reduced by $1/\sqrt{N}$, where N is the number of channels. Band-averaged polarization is given by $\langle P_o \rangle$. Similarly, the band-averaged fractional polarization is given by

$$\langle p_o \rangle = \frac{\langle P_o \rangle}{\langle I_{\text{pow}} \rangle}, \quad (2)$$

with fractional Q and U given by $u = U/I_{\text{pow}}$ and $q = Q/I_{\text{pow}}$. We compare band-averaged polarization values with previous studies. I_{pow} is the simple power-law fit to the measured peak total intensity:

$$I_{\text{pow}} = I_o \left(\frac{\nu}{\nu_o} \right)^\alpha. \quad (3)$$

The bias-correction method applied best approximates p_o when $p/\sigma_{qu} \gtrsim 3$. Here σ_{qu} is the rms noise in q and u . Table 2 summarizes the band-averaged total and polarized flux densities we derived for each source in each band, as well as the total intensity spectral index between the two bands.

The recent study by Jagannathan et al. (2017) characterized the effect of instrumental polarization leakage in interferometric imaging. Such effects are much reduced for long integrations with large coverage in parallactic angle and also for sources nearer to the pointing centre of the observed field. We use the secondary calibrators to calibrate on-axis leakages. We find leakage terms at the level of ~ 0.06 per cent. PKS 1934–638 was observed during each observing run as the primary calibrator. This source has no detectable polarization at L band and is a well-known Southern hemisphere flux calibrator (e.g. Reynolds 1994). Our average observed per cent polarization for PKS 1934–638 from the two frequency bands is $p = 0.06 \pm 0.07$ per cent, indistinguishable from the on-axis instrumental polarization. The most conservative estimate of instrumental polarization then, is to assume that PKS 1934–638 is unpolarized at this level of sensitivity and that the observed polarization is entirely instrumental. We have thus included in Table 2 an additional uncertainty of 0.06 per cent in quadrature to the formal noise derived error on the linear polarization degree.

The fractional polarization degree of each source in each band is listed in Table 3. The table also lists some polarization fractions at similar frequencies found in the literature from Eichendorf & Reinhardt (1979), who presented a collection of polarization properties (position angles, flux densities, and RMs) that were measured between 1965 and the middle of 1974 and these include data at 5.00,

Table 2. Band-averaged values of Stokes $\langle I \rangle_{\nu_c}$, band-averaged linear polarization intensities ($\langle P \rangle_{\nu_c}$ and bias-corrected $\langle P_o \rangle_{\nu_c}$), and also Q , U noise (σ_{QU}). The error in the averaged values are σ/\sqrt{N} , where N is the number of channels, σ is the standard deviation of the values, and ν_c is the frequency associated with the band. The low and high bands are indicated by subscripts/superscripts 1.35 and 1.85, respectively, indicating ν_c in GHz. All values of intensity have units of mJy per beam, abbreviated as mJy bm^{-1} . The last column displays the mean spectral index, $-\langle \alpha \rangle_{1.35}^{1.85}$, over the two KAT-7 bands.

Source	$\langle I \rangle_{1.35}$ (mJy bm^{-1})	$\langle P \rangle_{1.35}$ (mJy bm^{-1})	$\langle P_o \rangle_{1.35}$ (mJy bm^{-1})	$\langle \sigma \rangle_{QU}^{1.35}$ (mJy bm^{-1})	$\langle I \rangle_{1.85}$ (mJy bm^{-1})	$\langle P \rangle_{1.85}$ (mJy bm^{-1})	$\langle P_o \rangle_{1.85}$ (mJy bm^{-1})	$\langle \sigma \rangle_{QU}^{1.85}$ (mJy bm^{-1})	$\alpha_{1.85}^{1.35}$
NGC 1808	514.6 \pm 0.2	2.4 \pm 0.5	2.4 \pm 0.5	0.3	400.3 \pm 0.2	2.2 \pm 0.4	2.2 \pm 0.4	0.3	-0.82 \pm 0.01
NGC 1097	340.8 \pm 0.3	2.6 \pm 0.4	2.6 \pm 0.4	0.4	248.0 \pm 0.1	3.6 \pm 0.3	3.5 \pm 0.3	0.2	-1.06 \pm 0.02
J0240–231	5727.4 \pm 1.2	58.9 \pm 3.6	58.9 \pm 3.6	0.6	4861.1 \pm 0.8	101.5 \pm 3.1	101.5 \pm 3.1	0.5	-0.56 \pm 0.01
J0538–440	3923.0 \pm 0.3	27.6 \pm 2.8	27.6 \pm 2.8	0.6	3806.5 \pm 1.3	10.9 \pm 2.5	10.9 \pm 2.5	0.4	-0.10 \pm 0.01
PKS 1934–638	14876.9 \pm 0.5	–	–	1.5	13315.6 \pm 0.3	–	–	1.0	-0.36 \pm 0.01
PKS B0407–658	15130.7 \pm 1.0	18.7 \pm 9.6	18.7 \pm 9.6	1.5	10848.2 \pm 1.0	21.7 \pm 6.7	21.7 \pm 6.7	1.3	-1.12 \pm 0.01

Table 3. Degrees of linear polarization (in percentage) from Eichendorf & Reinhardt (1979) along with corresponding values from the KAT-7 bands in this work; $\langle p_o \rangle_{1.35}$ and $\langle p_o \rangle_{1.85}$. Subscripts are as in Table 2.

Source	$p_{5.00}$ (%)	$p_{2.73}$ (%)	$p_{1.67}$ (%)	$p_{1.43}$ (%)	$\langle p_o \rangle_{1.35}$ (%)	$\langle p_o \rangle_{1.85}$ (%)
NGC 1808	–	–	–	–	0.5 \pm 0.09	0.5 \pm 0.10
NGC 1097	–	–	–	–	0.8 \pm 0.11	1.4 \pm 0.11
PKS 1934–638	0.1	0.3	0.4	0.2	–	–
PKS B0407–658	0.5	0.4	0.6	0.4	0.2 \pm 0.07	0.2 \pm 0.07
J0240–231	5.4	3.4	1.7	1.8	1.0 \pm 0.07	2.1 \pm 0.07
J0538–440	–	–	–	–	0.7 \pm 0.08	0.3 \pm 0.07

2.73, 1.67, and 1.43 GHz – listed in Table 3. The typical observational errors in their data are in the range of 1–2 per cent in 1965, while being better at 0.5–1 per cent in 1974. Their data set contains observations for only three of our sources.

Linear polarization for just one of our targets, NGC 1097, was measured (Beck et al. 2002, 2005) at 22, 18, 6, and 3 cm (corresponding to ~ 1.36 , 1.66, 5.00, and ~ 10.00 GHz, respectively) and also at 4.80 GHz (Stil et al. 2009a). The integrated linear polarization degrees for NGC 1097 were reported at low significance where $p_{22\text{cm}} = 1.5 \pm 0.9$ per cent and $p_{18\text{cm}} = 1.3 \pm 1.0$ per cent at resolutions of 2–15 arcsec, but still consistent with our results calculated from the peak polarized flux: $p_{22\text{cm}} = 0.8 \pm 0.1$ per cent and $p_{16\text{cm}} = 1.4 \pm 0.1$ per cent (see Table 3). The source was undetected at 4.80 GHz where $p_o = 0.7 \pm 0.8$ per cent.

3.2 RM synthesis

The electric field vector of linearly polarized electromagnetic waves is rotated through the Faraday process, as the waves pass through a magnetoionic medium. The component of the magnetic field along our line of sight (LoS) can be probed through observations of this emission. Burn (1966) presented a technique that can resolve the different components of magnetic fields that alter the polarization of the incident radiation along the LoS. The technique was later expanded and dubbed ‘Rotation Measure Synthesis’ (RM Synthesis, for short) by Brentjens & de Bruyn (2005).

We apply the RM Synthesis technique to analyse the intrinsic polarization properties of our sources. The technique transforms $q(\lambda^2)$ and $u(\lambda^2)$ to $q(\phi)$ and $u(\phi)$ through a Fourier transform – ϕ is the Faraday depth, the LoS integral of the thermal electron number density (denoted as n_e) and magnetic field (\mathbf{B}) component along the observer’s LoS:

$$\phi = K_c \int_r^0 n_e \mathbf{B} \cdot d\mathbf{r}, \quad (4)$$

where K_c is a constant. Faraday rotation is quantified by the RM in units of rad m^{-2} . The RM is defined by the change in polarization angle $\chi(\lambda^2)$ as a function of the square of the wavelength, λ^2 :

$$\text{RM} = \frac{d\chi(\lambda^2)}{d\lambda^2}, \quad (5)$$

with the polarization angle given by

$$\chi = \frac{1}{2} \arctan \frac{u}{q}. \quad (6)$$

The Faraday depth and the RM are the same in the simplest case where there is only a single source of non-varying Faraday rotation. This is the so-called Faraday simple case (e.g. Brown et al. 2019), where polarization angle, χ , is related to the Faraday depth by

$$\chi = \chi_0 + \text{RM}\lambda^2. \quad (7)$$

Equation (7) may be valid at modest bandwidths, where $\chi(\lambda^2)$ is linear, but this is not a guarantee. In the case that the Faraday rotation plasma and the emitting plasma are mixed, this simple relationship cannot be assumed (Farnsworth, Rudnick & Brown 2011). Observations over broad bandwidths can display more complex Faraday rotation. To characterize such Faraday complexity we use the Faraday spectrum, denoted as $F(\phi)$ (Sun et al. 2015), referred to as the Faraday dispersion function (Brentjens & de Bruyn 2005). Such spectra can reveal Faraday components in polarized emission at different Faraday depths, ϕ . $F(\phi)$ is related to linear polarized intensity, P , by

$$P(\lambda^2) = \int_{-\infty}^{\infty} F(\phi) e^{2i\phi\lambda^2} d\phi. \quad (8)$$

To perform RM Synthesis, we used a grid of trial ϕ values, spaced by 0.1 rad m^{-2} ($\phi \in [-1000, 1000] \text{ rad m}^{-2}$), to measure $F(\phi)$. The inverse Fourier transform to derive $F(\phi)$ is approximated by a summation as observations do not cover a continuous wavelength

Table 4. Key parameters in our RM Synthesis analysis. f_c is the frequency range covered, ‘Target’ refers to the disc galaxy target’s observation run from which the quantities in the table are derived, $\delta\lambda^2$ is the spacing in λ^2 , λ_{\min}^2 and λ_{\max}^2 are the minimum and maximum values, respectively, of λ^2 corresponding to each band. $\phi_{\max\text{-scale}}$ is the largest Faraday depth scale to which our RM Synthesis analysis is sensitive and $|\phi_{\max}|$ is the largest depth we can detect. $\delta\phi'$ is the RMTF FWHM due to the non-continuous λ^2 coverage in this work, while $\delta\phi$ is the same for the continuous case.

f_c	Target	$\delta\lambda^2$ (m ²)	λ_{\min}^2 (m ²)	λ_{\max}^2 (m ²)	$\phi_{\max\text{-scale}}$ (rad m ⁻²)	$\delta\phi$ (rad m ⁻²)	$\delta\phi'$ (rad m ⁻²)	$ \phi_{\max} $ (rad m ⁻²)
1.251–1.962 GHz	NGC 1097	9.29×10^{-5}	2.33×10^{-2}	5.74×10^{-2}	134.6	53.1	101.0 ± 0.8	1.86×10^4
1.318–1.970 GHz	NGC 1808	9.18×10^{-5}	2.32×10^{-2}	5.17×10^{-2}	135.6	67.7	98.0 ± 0.8	1.89×10^4

range and also $\lambda > 0$ at all times:

$$F(\phi) \approx K \sum_k P(\lambda_k^2) e^{-2i\phi(\lambda_k^2 - \lambda_0^2)}, \quad (9)$$

where K is the inverse of the λ^2 integral of the sampling function, $W(\lambda^2)$ (which is zero everywhere beyond the limits of the observed λ^2 range but non-zero where the range is sampled):

$$K = \left(\int_{-\infty}^{\infty} W(\lambda^2) d\lambda^2 \right)^{-1}. \quad (10)$$

We progress through the following steps in applying RM Synthesis and deconvolving the intrinsic Faraday spectrum of each source.

(1) Calculate the Faraday spectrum for an input q, u array to avoid the effects of a varying spectral index across a broad-band that affects Q and U .

(2) Perform a deconvolution on the Faraday spectrum (i.e. perform an ‘RM CLEAN’; Heald, Braun & Edmonds 2009) with a loop gain of 0.1. This deconvolution is iterated in a computational loop set to terminate at the RM CLEAN threshold, explained below. The RM Synthesis analysis is done in the range -1000 to $+1000$ rad m⁻². RM ranges up to $|\phi_{\max}|$ (see Table 4) were also tested and no emission components were detected beyond ± 1000 rad m⁻².

(3) The result of the RM CLEAN loop is then a multicomponent model with each ϕ -component characterized by a position angle, peak amplitude of the Faraday spectrum ($|F(\phi)|$), and a ϕ value.

(4) We then create RM CLEANed $F(\phi)$, $q(\phi)$, $u(\phi)$, $q(\lambda^2)$, and also $u(\lambda^2)$ data sets based on the multicomponent model.

Table 4 lists the approximate values of parameters quantifying the limits of our KAT-7 data with regard to RM Synthesis. The predicted resolution in Faraday space, $\delta\phi$, is slightly worse (67.7 rad m⁻²) in the observations of NGC 1808, J0538–440, PKS B0407–658, and PKS 1934–638 than for NGC 1097 and J0240–231 (where $\delta\phi = 53.1$ rad m⁻²) due to a larger fraction of the band being flagged in the former case. The λ^2 sampling of KAT-7 allows detection of RM components up to $|\phi_{\max}| \sim 1.86\text{--}1.89 \times 10^4$ rad m⁻², where $|\phi_{\max}| \approx \sqrt{3}/\delta\lambda^2$ (Brentjens & de Bruyn 2005). $|\phi_{\max}|$ is calculated at the 3.9 MHz channel width where $\delta\lambda^2 \sim 9 \times 10^{-5}$ m² as listed in Table 4. The worst-case unaccounted rotation (occurring at the lowest wavelength where $\lambda_{\max}^2 \approx 0.05$ m²) would therefore be $\sim 196^\circ$. Note that the values listed in Table 4 are strictly correct for only continuous frequency coverage between the high and low frequencies of the observations. The effect of the absence of the central band is to create strong ‘sidelobes’ of the ‘dirty’ Faraday transfer function. The quantity $\delta\phi'$ lists the measured full width at half-maximum (FWHM) of the main lobe of the ‘dirty’ RM transfer function (RMTF), and represents the actual Faraday resolution of our RM Synthesis spectra.

We quantify the statistical significance of ϕ -component detections by imposing an amplitude threshold in the RM CLEAN algorithm. Our target confidence is set at $\gtrsim 99$ per cent. To this end, we select 1000 off-source points in each of our image cubes, extract the Q , U , and

I fluxes in a similar fashion as described previously. We then apply our RM CLEAN algorithm and for each science target image cube, we determine the maximum amplitude $|F|_{\max}$ from the set of off-source positions, above which the Probability($|F|_{\max, i} \geq |F|_{\max}$) ≤ 0.001 . This gives us a 0.999 confidence of individual RM components for each of our science targets. Finally, we combine this limit with the instrumental polarization leakage error estimate of 0.06 per cent by adding them in quadrature and using the result as the lower threshold for detection of Faraday components. Table 5 lists the thresholds and detected Faraday components for each source. We compared the RM Synthesis results with and without a weighting of σ_{qu}^{-2} applied and found that the results do not differ. We thus apply no weighting in our final RM Synthesis analysis.

We characterize Faraday complexity according to the second moment about the mean, σ_ϕ , of the RM CLEAN model spectrum (Brown 2011; Anderson et al. 2015) given by

$$\sigma_\phi = \sqrt{\frac{\sum_i (\phi_i - \mu_\phi)^2 |F(\phi_i)|}{\sum_i |F(\phi_i)|}}, \quad (11)$$

with

$$\mu_\phi = \frac{\sum_i \phi_i |F(\phi_i)|}{\sum_i |F(\phi_i)|}. \quad (12)$$

A Faraday simple spectrum will have $\sigma_\phi \approx 0$ rad m⁻², while the complex counterpart will have σ_ϕ very different from zero. We could not reliably resolve the complexity in individual components given that $\delta\phi' \sim 100$ rad m⁻² but we classify a source spectrum as Faraday complex (denoted by C in Table 5) or simple (denoted as S) according to σ_ϕ calculated from all components that are detected above the noise. Table 5 summarizes the ϕ -components we found near the RM CLEAN threshold level for each source. The *major/main/prominent* (highest signal-to-noise ratio, S/N) components show some width manifesting as a number of *minor* (low S/N) components separated by less than $\delta\phi'$ from the main component. A Faraday emission component, with $\phi = \phi_i$, is called *main/most prominent* if $|F(\phi_i)| = \text{maximum}(|F(\phi)|)$, and *minor* otherwise. The ϕ_i of minor components within $\delta\phi'$ of each other is combined to give μ_ϕ , of detections as listed in Table 5. The 1σ deviations of these minor components are reported as the error on μ_ϕ .

Figs 2(a)–(c) show the λ^2 dependence of q , u , and p_o along with the model spectra reconstructed from the RM Synthesis CLEAN components for each source. Figs 3(a)–(f) show the ‘dirty’ $F(\phi)$ for each source along with the RM Synthesis CLEAN components. The strong sidelobes of the RM transfer function from the missing mid-band data set are visible in the dirty spectra. A ‘CLEANed’ $F(\phi)$ is also plotted, constructed by convolving the RM CLEAN components with an idealized window function corresponding to full frequency coverage over the frequency range of the observations. The results for each source are discussed below.

J0240–231. This is the most polarized source in our sample with several Faraday components that are above the detection threshold –

Table 5. RM CLEAN components for each of our sources. $|F|_{\text{comp}}$ is the peak of the Faraday spectrum derived during the RM CLEAN algorithm, μ_ϕ (equation 12) represents the Faraday depth of each component, σ_ϕ is the second moment of the RM CLEAN model spectrum (Brown 2011; Anderson et al. 2015), and the RM CLEAN threshold. The column ‘Class’ shows classification of each source as Faraday complex (C) or Faraday simple (S) according to the value of σ_ϕ . The last three columns show galactic coordinates (l , b) and also estimate of the smooth Galactic foreground Faraday depths, ϕ_{gal} , at the locations of each our targets (Oppermann et al. 2012, 2015).

Source	$ F _{\text{comp}}$ (%)	μ_ϕ (rad m ⁻²)	σ_ϕ (rad m ⁻²)	Threshold (%)	Class	l (°)	b (°)	ϕ_{gal} (rad m ⁻²)
J0240–231	1.62	8.0 ± 0.1	37.7	0.061	C	39.415	−23.224	9.1
	1.48	33.9 ± 1.4						
	1.29	−38.1 ± 4.8						
	0.69	−110.8 ± 2.0						
J0538–440	0.45	91.2 ± 9.6	40.8	0.061	C	84.491	−43.997	42.8
	0.17	9.6 ± 1.4						
NGC 1097	1.09	−21.5 ± 8.9	120.2	0.072	C	41.579	−30.275	9.1
	0.81	44.6 ± 3.5						
	0.47	147.3 ± 1.9						
	0.23	−520.5 ± 0.3						
	0.23	793.0 ± 0.3						
	0.22	619.9 ± 0.8						
NGC 1808	0.43	−0.5 ± 1.3	49.4	0.065	C	76.926	−37.513	35.3
	0.24	144.7 ± 0.7						
	0.12	−291.6 ± 0.1						
PKS 1934–638	–	–	–	0.061	–	65.146	−63.713	44.7
PKS B0407–658	0.15	8.0 ± 0.1	0.0	0.061	S	62.085	−65.753	23.9

see Fig. 3(a). The most prominent component, as listed in Table 5, is roughly consistent with the value of 14.4 ± 2.2 rad m⁻² measured by Taylor et al. (2009) based on a simple two-point measure of polarization angle. Several Faraday components are seen in Fig. 3(a) and indicate that the fainter components are Faraday thick. Generally, the Faraday spectrum is complex with $\sigma_\phi = 37.7$ rad m⁻².

J0538–440. This source shows two relatively thick components, each composed of a number of closely adjacent minor components (Fig. 3b). The peaks secondary in prominence to the two most prominent, seen as the lower level CLEAN components in Fig. 3(b), have been assumed to form part of the thick components. The source is classified as Faraday complex with $\sigma_\phi = 40.8$ rad m⁻².

NGC 1097. Displays the most complex Faraday depth profile amongst the two disc galaxy sources – Fig. 3(c), with thick components at low ϕ and some faint components just above the threshold at large $|\phi| > 500$ rad m⁻² (particularly near −520, 590, and 790 rad m⁻²). The most prominent component seems Faraday thick, with minor components spanning the range $\phi = -11$ to -32 rad m⁻². This source is classified as Faraday complex and displays the largest dispersion in Faraday space with $\sigma_\phi = 120.2$ rad m⁻². The component detected at 147.3 ± 1.9 rad m⁻² most likely has some sidelobe emission from the first sidelobe, at $\phi = +129.0$ rad m⁻², but cannot be solely due to the sidelobe as the same location reflected about 0.0 rad m⁻² does not show the same detectable signal.

NGC 1808. Displays a Faraday spectrum with three Faraday-thin components – Fig. 3(d). We detect three prominent ϕ -components. The component with $\mu_\phi = 144.7 \pm 0.7$ rad m⁻² is close to the first RMTF sidelobe located at $\phi = +140.5$ rad m⁻², however $|F(\phi)|$ is more than $3 \times$ the detection threshold, and its presence is clearly seen as an asymmetry in the dirty Faraday spectrum. The component located at $\mu_\phi = -291.6 \pm 0.1$ rad m⁻² also appears to coincide with the second RMTF sidelobe. The dirty Faraday spectrum hints at a component near the positive (i.e the RMTF in the $\phi > 0.0$ rad m⁻² region) second sidelobe but there are no detectable CLEAN

components in the region. The source is Faraday complex with $\sigma_\phi = 49.4$ rad m⁻².

PKS 1934–638. This source is undetected in polarization. The peak emission in the Faraday spectrum is near $\phi = 50.0$ rad m⁻² with a peak $|F| \approx 0.061$ per cent (≈ 0.064 per cent when the residuals are added in), which is set as the instrumental polarization level. Fig. 3(e) displays the Faraday spectrum.

PKS B0407–658. The Faraday spectrum of this source displays a single Faraday-thin component and does not show any other components above the threshold (Fig. 3f). The lone component is detected at $\mu_\phi = 8.0 \pm 0.1$ rad m⁻² with a peak $|F| = 0.15$ per cent.

3.2.1 Galactic foreground and ionospheric effects

Foreground RM contributions, such as from the ISM in the Galaxy and the Earth’s ionosphere, can also be significant along the LoS. This is especially so at frequencies below 1 GHz (e.g. Erickson et al. 2001). However, most of the Galactic foreground emission is resolved out by an interferometer because of missing short spacings. There can be small-scale structure in Stokes Q and U that the interferometer detects, and that can in principle become visible as a component in the Faraday depth spectrum. These are sometimes called Faraday ghosts and are independent of position on the sky (e.g. Shukurov & Berkhuijsen 2003).

Oppermann et al. (2012, 2015) have created a map of the smoothly varying Galactic Faraday component, ϕ_{gal} , by utilizing all-sky Faraday RMs of compact background sources from the literature. We list ϕ_{gal} at the locations of each of our targets in Table 5. The effect of the Galactic ISM on an extragalactic source is mainly as a foreground Faraday-rotating screen. The correction for this effect is to subtract ϕ_{gal} from detected components. In this work, we note that this effect does exist and has possibly affected our measurements but we do not correct for it as it is not adequately constrained.

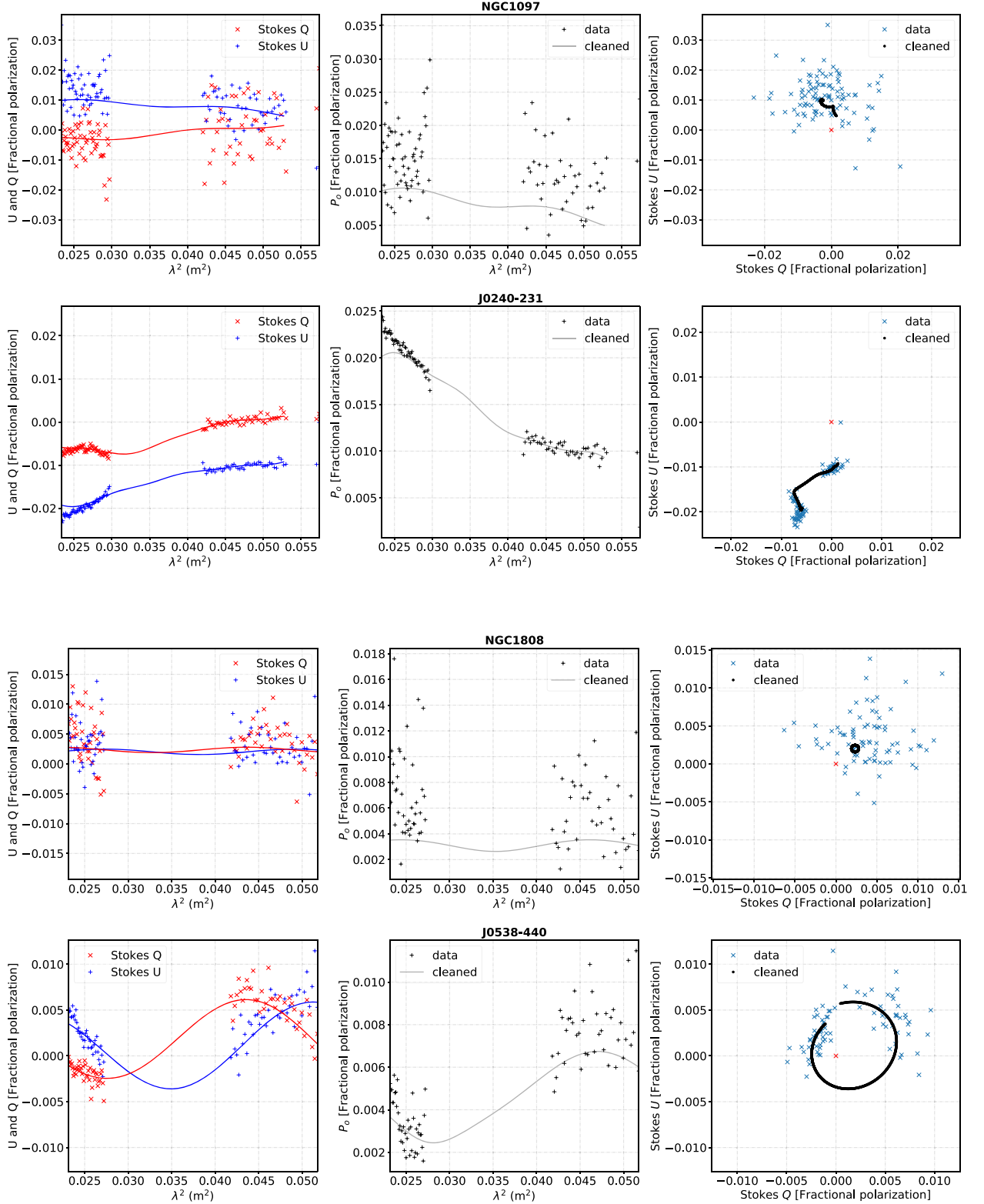
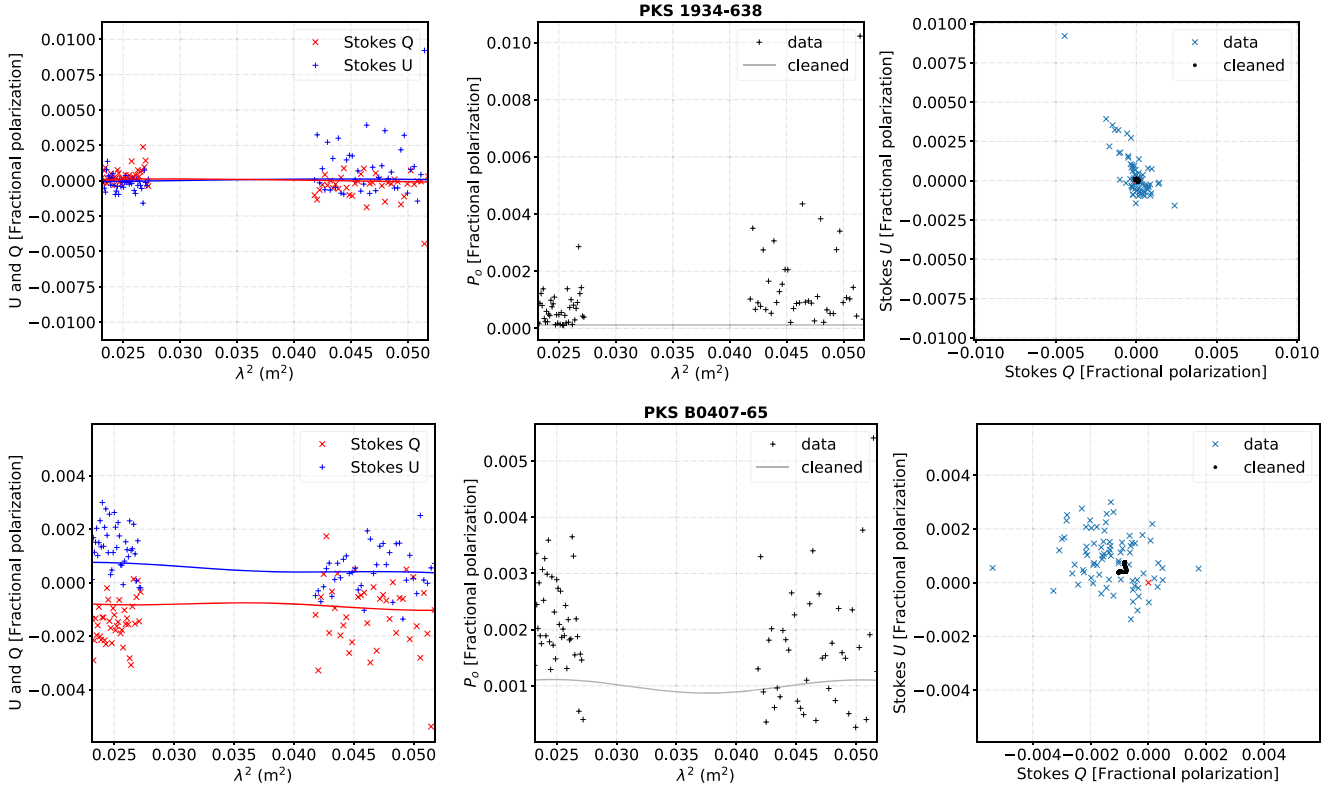


Figure 2. Polarization measurements, fractional Stokes Q and U (with errors, error bars omitted, of the order $\sim 10^{-5}$), and also total linear polarization p_0 , across the KAT-7 band along with the RM ‘CLEANED’ versions of each quantity – in each panel, shown as curves in the leading two plots and black points in the trailing plot.

In an effort to analyse Faraday complexity introduced by the ionosphere at the different observation epochs, we look for rotation in the polarization angle (χ) as a function of time. This we do by

producing full Stokes images for each scan of the phase calibrator for each epoch. We then restrict our analysis to those scans that have Stokes V flux and p_0 that are below the leakage, and also Stokes I that

Figure 2. *continued.*

is above 3σ . We then analyse the resulting q , u , and χ as functions of scan, i.e. time. However, the resulting images prove to be too noisy to lead to a conclusion on ionospheric effects.

Typical Faraday RM variations from the ionosphere are up to ~ 1 rad m^{-2} . Typical variations are of order 1 rad at 1 m wavelength as $\chi(\lambda^2) \sim \lambda^2$. Variations of this order are seen at 327 MHz at the VLA (Erickson et al. 2001). At our lowest frequency even a 90° change at 327 MHz scales to only about 6° at 1300 MHz, and 2.5° at 1900 MHz. This is small enough to ignore. Note also that it is only the time variation during an epoch that will affect the observations. A constant offset from the ionosphere will be accounted for by the calibration. In the absence of an absolute polarization calibrator, the solution is referenced to the mean feed position angle that is not expected to change as it is hardwired into the instrument. We therefore conclude that the expected ionospheric rotation is negligible and thus not corrected.

4 DISCUSSION

Our RM Synthesis analysis shows Faraday complexity, of varying degrees, in four of the six sources in our KAT-7 sample. The most polarized source, J0240–231, displays the most complex Faraday spectrum with several emission components (see Fig. 3a and Table 5). The presence of Faraday thick components can be seen in the Faraday spectrum as well. Our analysis clearly resolves this complex spectrum both near the main peak and the smaller peaks. This source also shows the most rapidly changing polarized intensity that depolarizes toward larger wavelengths (see the bottom panel of Fig. 2a). The behaviour of $q(\lambda^2)$ and $u(\lambda^2)$ is also far from the simple sinusoidal behaviour, displayed by a single Faraday component, and this deviation is best illustrated in Fig. 2(a). This plot resembles

a circle of constant radius in the Faraday simple case, but we see quite a substantial deviation from this. Our RM CLEAN model can however reproduce the behaviour we see in the spectra. The $q(\lambda^2)$ and $u(\lambda^2)$ profiles have little scatter and are in good agreement with our RM CLEAN model adding confidence to the derived ϕ -components. Anderson et al. (2016) show that Faraday complexity is common in highly polarized sources and our observation in the case of J0240–231 is in accord. Schnitzeler et al. (2019), however, fit one Faraday component to the J0240–231 spectrum through a QU -fitting analysis. The source was found to be significantly polarized at a reference frequency of 2.1 GHz with $p = 2.28$ per cent suggesting agreement with the measurement made in this work at the high, 1.85 GHz, band.

Faraday complex sources are also found to more likely have steep spectra in the optical and radio (e.g. Scarpa & Falomo 1997; Anderson et al. 2015). We find J0240–231 to have only a mildly steep total intensity spectrum of $\langle \alpha \rangle_{1.85}^{1.35} = -0.562 \pm 0.006$. Our results for J0240–231 agree with previous findings that the source is significantly polarized with $p_{4.9\text{GHz}} = 3.47$ per cent (Perley 1982) and $p_{1.4\text{GHz}} = 0.8$ per cent (Taylor et al. 2009). The object is a bright compact radio source (Fanti et al. 1990, 2000), however, the double-lobed structure reported by Dallacasa et al. (1998) may contribute to the observed Faraday complexity as radio lobes can be sites of complex Faraday structures (e.g. Banfield et al. 2019; Sebokolodi 2019).

The case of J0538–440 is another where we find evidence of Faraday complexity, although not to the level of J0240–231 (contrast Fig. 3b with Fig. 3a). The behaviour of $q(\lambda^2)$ and $u(\lambda^2)$ profiles for this source (bottom panels of Fig. 2b) resembles that of a simple one-component Faraday spectrum with somewhat larger scatter than in J0240–231, indicating that perhaps higher sensitivity observations

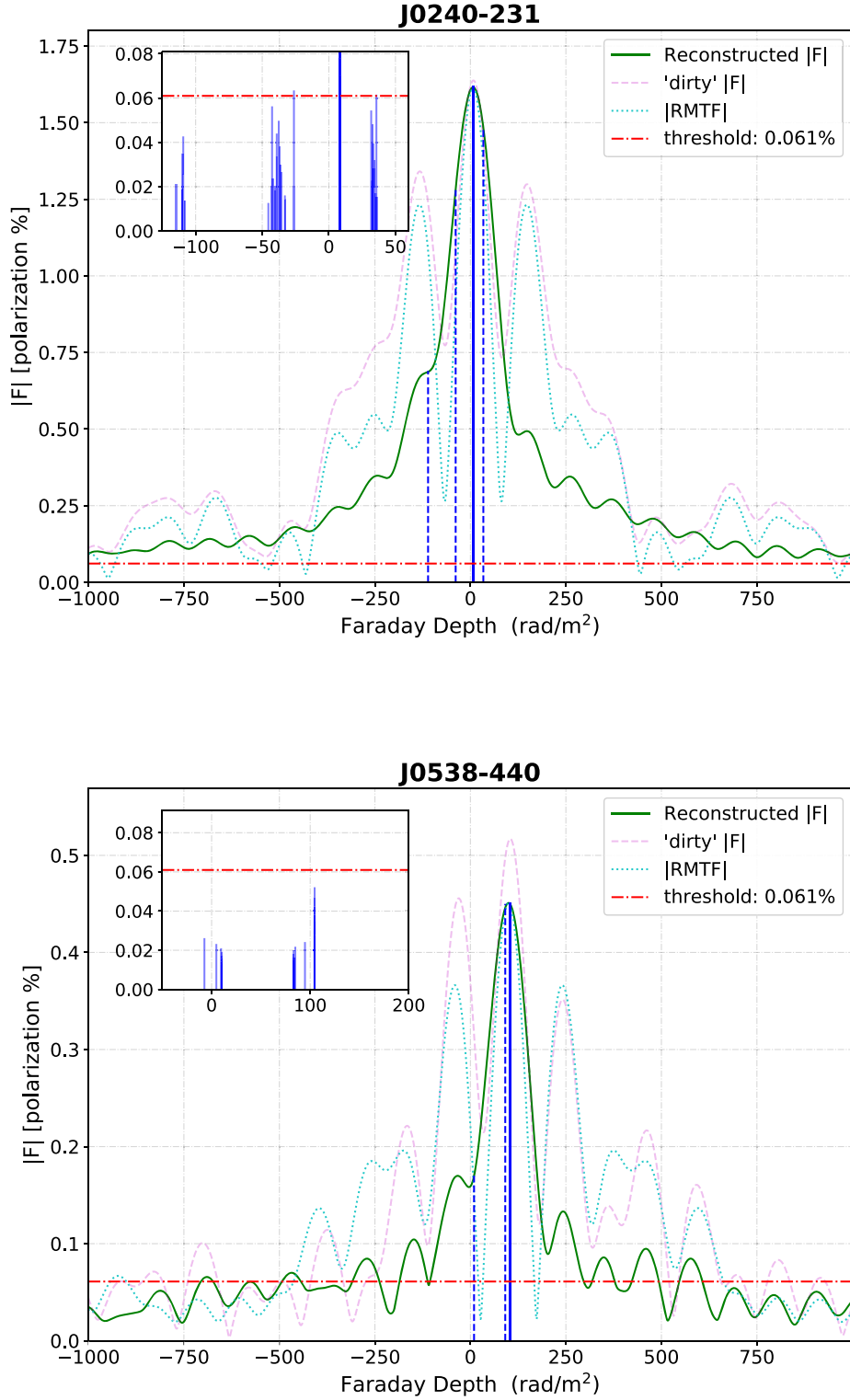
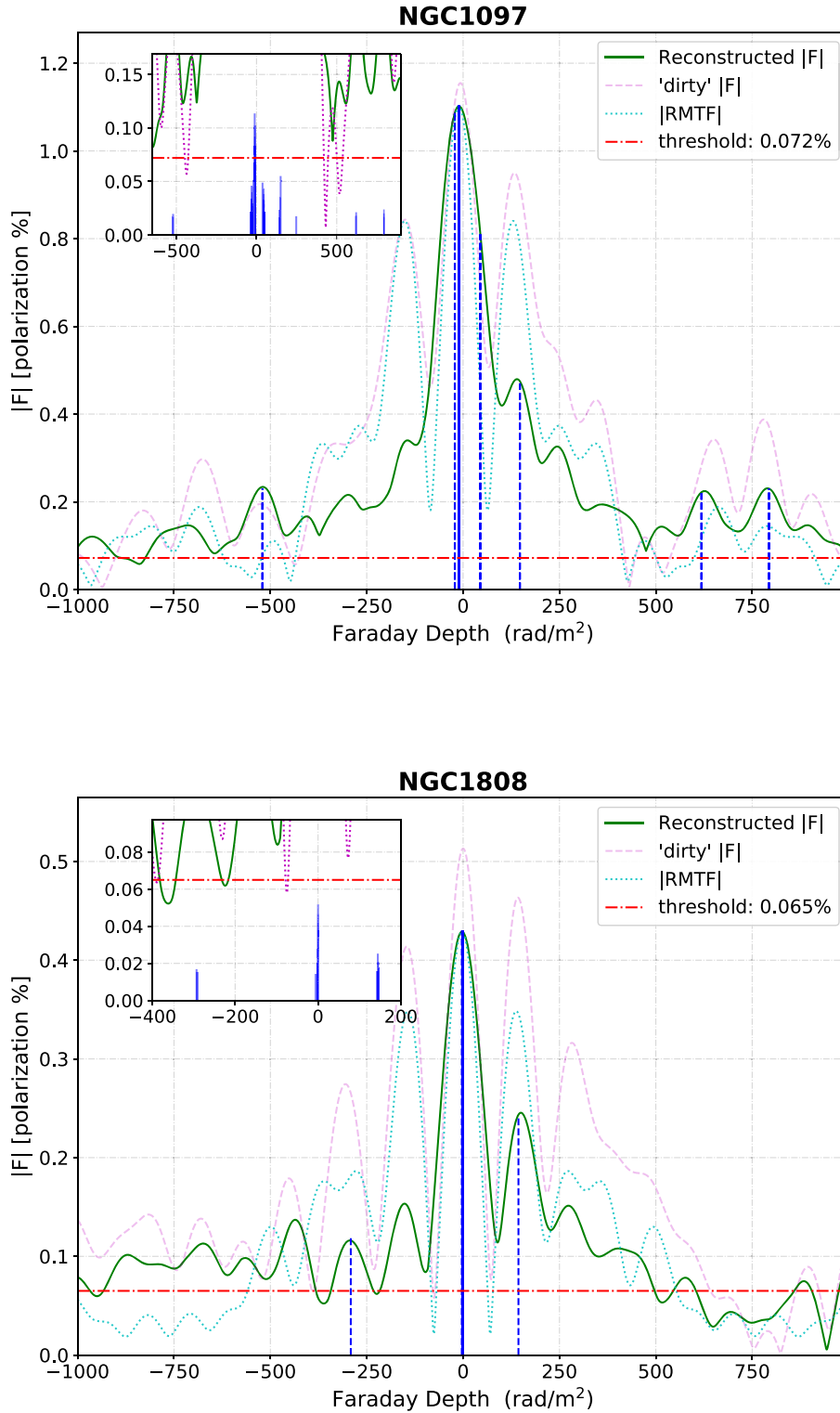
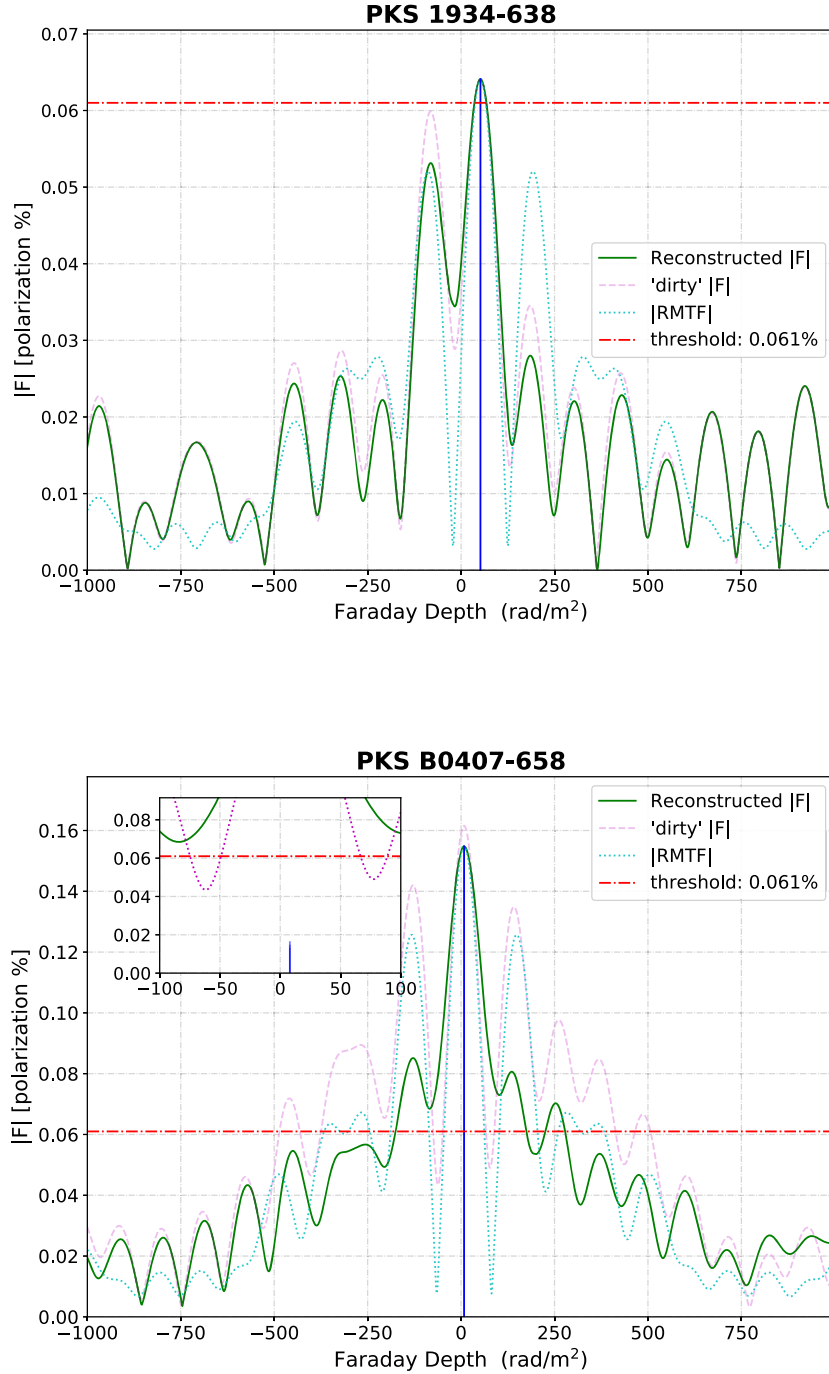


Figure 3. The CLEAN Faraday spectrum, $|F(\phi)|$, resulting from our RM CLEAN algorithm (green solid curve) with residuals added back in. Top left: zoom-in of the main plot showing deconvolved Faraday emission components (blue solid vertical lines show the peak $\gamma|F(\phi_i)|$ at the corresponding ϕ_i of each CLEAN component in the vicinity of the brightest components) – γ is the loop gain parameter set to 0.1. The red dash-dotted horizontal line is the RM CLEAN threshold as quoted in the legend. The dashed magenta curve is the 'dirty' Faraday spectrum, while the dotted cyan curve is the RMTF scaled to the amplitude of the deconvolved $|F(\phi)|$ and with peak located at the Faraday depth corresponding to the peak of $|F(\phi)|$. Blue vertical dashed lines in the unzoomed plot display the value of μ_ϕ of each detected Faraday component as quoted in Table 5.

Figure 3. *continued.*

with a larger λ^2 coverage and thus Faraday depth resolution are needed to further explore the behaviour and further resolve the ϕ -components. This source is highly polarized, ~ 18 per cent, in the optical (Giommi, Ansari & Micol 1995; Scarpa & Falomo 1997) and has possibly suffered significant depolarization at L band. Our RM CLEAN implementation is able to model the behaviour well at the $\sim 2.5\sigma$ level imposed. The J0538–440 total intensity spectrum is

more time variable but appears flatter than in the case of J0240–231 with $\langle \alpha \rangle_{1.85}^{1.35} = -0.098 \pm 0.003$ and a lower degree of polarization of less than 1 per cent at L band. This is another source that is also in the Schnitzeler et al. (2019) catalogue where it is found to be Faraday simple and highly polarized, $p = 3.32$ per cent (at RM $\approx 60 \text{ rad m}^{-2}$), at 2.1 GHz in contrast with $p = 0.3 \pm 0.07$ per cent measured at the high band in this work. The total intensity of this


 Figure 3. *continued.*

source is a highly time variable and we suspect that the significant discrepancy we observe is due to this variability as sources that are variable in total intensity may also be variable in polarization (Anderson et al. 2019). We observed this source in late 2016 July as compared to the 2011 observations of Schnitzeler et al. (2019). The authors obtained their results from fitting simple Faraday components to the sources spectrum that differs significantly from our open ended approach in which our model can contain as many (CLEAN) components as is necessary to fit the data while minimizing errors. We do not apply constraints to this work. Another contributing factor to observed discrepancies is the difference in frequency bands, L and

S bands. At L band, it is typical that polarization is lower than at S band, especially for compact sources.

NGC 1097 displays the most complex Faraday spectrum in the two disc galaxies. The Faraday spectrum of the source shows the largest spread in ϕ (see Fig. 3c). Low S/N emission at comparatively large ϕ could be contaminated by sidelobes of the RMTF but may also be due to small regions of highly magnetized media that may emit at large Faraday depths. These small regions may coincide with highly turbulent regions near the core of a galaxy (e.g. Pasetto et al. 2018). The low S/N of these components may also point to depolarization that may be due to turbulence of the magnetized plasma in the bar and circumnuclear ring regions of the galaxy and/or frequency-

dependent depolarization that diminishes the polarized signal at L band. Higher resolution measurements by Beck et al. (2005) do not show evidence of high rotation-measure regions in NGC 1097 (the authors report ϕ in the range -200 to $+200$ rad m $^{-2}$ from 50 MHz bandwidth observations at 4.86 and 8.46 GHz) and so we may be detecting a noise contribution to the signal that is unaccounted for. More sensitive observations with a wider bandwidth coverage will be required to confirm these faint components. Polarimetry at higher frequency may also be able reveal such components in the absence of internal depolarization effects (e.g. Anderson et al. 2016).

NGC 1808 is also classified as Faraday complex in our RM CLEAN analysis with the detection of three Faraday-thin components – see Fig. 3(d). The complexity classification in this disc galaxy is due to the fairly large spread of Faraday depths found. Our RM CLEAN implementation shows subtle structure of the polarized spectra of the galaxy, while the data show a low S/N and significant scatter at the wavelengths observed. A larger λ^2 coverage with higher sensitivity is required as we will be able to better constrain the polarized spectra of such Faraday complex and low polarized S/N sources (e.g. O’Sullivan et al. 2012; Anderson et al. 2015, 2016).

The apparent spectral indices we derive suggest agreement with those expected for spiral galaxies, $\alpha = -0.8$ to -1.0 (Beck 2015). We find $\langle\alpha\rangle_{1.85}^{1.35} = -0.815 \pm 0.011$ for NGC 1808 and $\langle\alpha\rangle_{1.85}^{1.35} = -1.060 \pm 0.020$ for NGC 1097. NGC 1808 is weakly polarized, with $p < 1$ per cent in accordance with models of the integrated polarization of galaxies at 1.4 GHz (Stil et al. 2009b; Sun & Reich 2012). The source also shows evidence of possible polarized emission components emitted at high Faraday depths as suggested by low level (undetected) large ϕ emission in the deconvolved Faraday spectrum.

PKS 1934–638 is not detected in polarization. This source shows polarization levels that are consistent with noise across the L band. Eichendorf & Reinhardt (1979) quote polarization levels that are of order double what we have found with observational errors typically being in the range of 1–2 per cent in 1965 and lowering to 0.5–1 per cent in 1974. We interpret their result as upper limits. Our results are comparable to the earlier data (Table 3) showing that PKS 1934–638 is undetected in polarization.

PKS B0407–658, however, is found to have one detected component – Fig. 3(f). The $q(\lambda^2)$, $u(\lambda^2)$, and $p(\lambda^2)$ profiles of this source are not as featureless as in PKS 1934–638 and may require a larger λ^2 coverage to further resolve the Faraday spectrum. More sensitive MeerKAT commissioning observations show that PKS B0407–658 is polarized at ~ 0.03 – 0.05 per cent levels in the L band (Hugo et al. 2018).

There exists no known relation between Faraday complexity and morphological complexity at the angular scales probed here (e.g. Anderson et al. 2015), so even the most complex sources may be much simpler in morphology. More sensitive and higher resolution broad-band spectropolarimetry studies are needed to further resolve Faraday components in Faraday depth, across wide bandwidths, at higher frequencies, and also spatially with high S/N as will be the case with the next generation of radio telescopes such as MeerKAT.

5 SUMMARY AND CONCLUSIONS

We have explored the properties of broad-band linear polarization in two galaxies and four AGN-powered sources in direct pointing mode with the KAT-7 array. The observations were made at the KAT-7 low band ($\nu_c \sim 1350$ MHz) and high band ($\nu_c \sim 1850$ MHz) with the mid band ($\nu_c \sim 1550$ MHz) discarded due to RFI. The synthesized beam at the low band was ~ 4.9 arcmin and the same at the high band was

~ 3.9 arcmin. We detect linear polarization down to ~ 2 mJy bm $^{-1}$. We employed the open-ended method of RM Synthesis and RM CLEAN (Burn 1966; Brentjens & de Bruyn 2005; Heald et al. 2009) that enabled us to resolve the Faraday depth components down to a Faraday depth resolution $\delta\phi' \sim 100$ rad m $^{-2}$ along our LoS. Our conclusions are summarized as follows.

(i) The total intensity spectra that we were able to measure are in general agreement with previous studies.

(ii) Our spectropolarimetric analysis shows Faraday simple, PKS B0407–658, unpolarized, PKS 1934–638, and Faraday complex sources, J0538–440, J0240–231, NGC 1097, and NGC 1808. The most polarized objects display more complex Faraday spectra that are indicative of multiple emitting and/or diminishing components at different Faraday depths along the LoS.

(iii) Our RM Synthesis analysis did not detect Faraday complexity in the least polarized source (PKS B0407–658) but such sources may still have some level of underlying Faraday complexity that is beyond detection. The two disc galaxy sources are classified as complex and have been found to have emission components over larger Faraday depth ranges – suggesting complex magnetoionic emission environments along the LoS.

(iv) The BL Lac source, J0538–440, and QSO, J0240–231, are the most highly polarized in the literature and most Faraday complex of our direct pointing sources. Our measurements support earlier findings that Faraday complexity is expected in most polarized sources. The lower Faraday complexity found in the other less polarized sources may require more sensitive wide band and thus higher ϕ resolution studies to further tease out the properties that may be beyond our current data. Low-polarization sources may also have a significant degree of spectropolarimetric complexity that is beyond detection in studies with similar sensitivity and resolution limitations as ours.

(v) We determine the on-axis instrumental polarization level in KAT-7 to have an upper limit of 0.06 per cent. This is based on the polarization level measured for the hitherto unpolarized GPS source PKS 1934–638 spectrum.

(vi) Some BL Lac sources, such as J0538–440, are particularly variable with time suggesting more complexity that should be accounted for when such sources are used as calibrators especially over broad bandwidths.

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6 DATA AVAILABILITY

The data underlying this paper can be shared upon reasonable request to the corresponding author.

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